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Tsallis Distribution as a Λ-Deformation of the Maxwell-Jüttner Distribution

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Abstract: Currently, there is no widely accepted consensus regarding a consistent thermodynamic framework within the special relativity paradigm. However, by postulating that the inverse temperature 4-vector, denoted as $\boldsymbol{\beta}$, is future-directed and timelike, intriguing insights emerge. Specifically, it is demonstrated that the q-dependent Tsallis distribution can be conceptualized as a de-Sitterian deformation of the relativistic Maxwell-Jüttner distribution. In this context, the curvature of the de Sitter space-time is characterized by $\sqrt{\lambda/3}$, where λ represents the cosmological constant within the λCDM standard model for Cosmology. For a simple gas composed of particles with proper mass m, and within the framework of quantum statistical de Sitterian considerations, the Tsallis parameter q exhibits a dependence on the cosmological constant given by $q = 1 + \ell_c\sqrt{\lambda}/\mathfrak{n}$ where $\ell_c = \hbar/mc$ is the Compton length of the particle and $\mathfrak{n}$ is a positive numerical factor. This formulation establishes a novel connection between the Tsallis distribution, quantum statistics, and the cosmological constant, shedding light on the intricate interplay between relativistic thermodynamics and fundamental cosmological parameters.

Keywords: Maxwell-Jüttner distribution; Tsallis distribution; de Sitter quantum field; ΛCDM standard model

MSC: 82B30;83A05;81R10;81T20

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1. Preamble: Temperature, Heat, Entropy, that Obscure Objects of Desire

It is opportune to start out this contribution by quoting what de Broglie was writing in [1] about the relation between entropy invariance and relativistic variance of temperature (translated from french):

> *It is well known that entropy, alongside the spacetime interval, electric charge, and mechanical action, is one of the fundamental "invariants" of the theory of relativity. To convince oneself of this, it is enough to recall that, according to Boltzmann, the entropy of a macroscopic state is proportional to the logarithm of the number of microstates that realize that state. To strengthen this reasoning, one can argue that, on the one hand, the definition of entropy involves a integer number of microstates, and, on the other hand, the transformation of entropy during a Galilean reference frame change must be expressed as a continuous function of the relative velocity of the reference frames. Consequently, this continuous function is necessarily constant and equal to unity, which means that entropy is constant.*

Let us now give more insights about what "relativistic thermodynamics" could be. In relativistic thermodynamics (i.e., in accordance with special relativity) there exist three points of view [2], distinguished from the way heat ΔQ and temperature T transform under a Lorentz boost from frame $\mathcal{R}_0$ (e.g., laboratory) to comoving frame $\mathcal{R}$ with velocity $\mathbf{v} = v\hat{\mathbf{n}}$ relative to $\mathcal{R}_0$ and Lorentz factor

$$\gamma(v) = \frac{1}{\sqrt{1 - v^2/c^2}}\,. \tag{1}$$

(a) the covariant viewpoint (Einstein [3], Planck [4], de Broglie [1] ...),

$$\Delta Q = \Delta Q_0 \gamma^{-1}\,, \quad T = T_0 \gamma^{-1}\,. \tag{2}$$

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(**b**) the anti-covariant one (Ott [5], Arzelies [6], ...),

$$\Delta Q = \Delta Q_0 \gamma\,, \quad T = T_0 \gamma\,. \tag{3}$$

(**c**) the invariant one, "nothing changes" (Landsberg [7,8], ...),

$$\Delta Q = \Delta Q_0\,, \quad T = T_0\,. \tag{4}$$

Also note that for some authors (Landsberg [9], Sewell [10], ...) "there is no meaningful law of temperature under boosts".

In this paper we adopt the viewpoint Section **a** and review de Broglie's arguments in 2. In Section 3 we remind the construction of the so-called Maxwell-Jütner distribution presented by Synge in []. We then present in Section 4 the de Sitter space-time, its geometric description as a hyperboloid embedded in the $1+4$ Minkowski space-time, and give some insights of the fully covariant quantum field theory of free scalar massive elementary systems propagating on this manifold. We then develop in Section 5 our arguments in favor a a novel connection between the Tsallis distribution, quantum statistics, and the cosmological constant, shedding light on the intricate interplay between relativistic thermodynamics and fundamental cosmological parameters. A few comments end our paper in Section 6.

2. Relativistic Covariance of Temperature According to de Broglie (1948)

We here give an account of the de Broglie arguments given in [1] in favour of the covariant viewpoint (a).

Let us consider a body $\mathcal{B}$ with proper frame $\mathcal{R}_0$, and total proper mass M_0. It is assumed to be in thermodynamical equilibrium with temperature T_0 and fixed volume V_0 (e.g., a gas enclosed with surrounding rigid wall). Let us then observe $\mathcal{B}$ from an inertial frame $\mathcal{R}$ in which $\mathcal{B}$ has constant velocity $\mathbf{v} = v\hat{\mathbf{n}}$ relative to $\mathcal{R}_0$. We suppose that a source in $\mathcal{R}$ provides $\mathcal{B}$ with heat ΔQ. In order to keep the velocity $\mathbf{v}$ of $\mathcal{B}$ constant a work W has to be done on $\mathcal{B}$. Its proper mass is consequently modified $M_0 \to M_0'$. Then, from energy conservation,

$$(M_0' - M_0)\gamma c^2 = \Delta Q + W\,, \quad \gamma = \gamma(v) = \frac{1}{\sqrt{1 - v^2/c^2}}\,, \tag{5}$$

and the relativistic 2nd Newton law,

$$\Delta P = M_0'\gamma v - M_0 \gamma v = \int F \mathrm{d}t = \frac{1}{v}\int F v \mathrm{d}t = \frac{W}{v}\,, \tag{6}$$

we derive

$$\Delta Q = \frac{c^2}{v^2}\gamma^{-2} W = (M_0' - M_0)c^2\gamma^{-2}\,. \tag{7}$$

In the frame $\mathcal{R}_0$ there is no work done (the volume is constant), there is just transmitted heat $\Delta Q_0 = (M_0' - M_0)c^2$. By comparison with (7) one infers that heat transforms as

$$\Delta Q = \Delta Q_0 \gamma^{-1}\,. \tag{8}$$

Since the entropy $S = \int \frac{\mathrm{d}Q}{T}$ is relativistic invariant, $S = S_0$, temperature finally transforms as

$$T = T_0 \gamma^{-1} \tag{9}$$

3. Maxwell-Jüttner Distribution

We now present a relativistic version of the Maxwell-Bolztman distribution for simple gases, namely Maxwell-Jüttner distribution. We follow the derivation given by Synge in [14], see also [15],

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and the recent [16] for a comprehensive list of references. Note that this distribution is defined on the mass hyperboloid and not expressed in terms of velocities (see the recent [17] and references therein).

Our notations [18] for event 4-vector $\underline{x}$ in the Minkowskian space-time $\mathbb{M}_{1,3}$ and for 4-momentum $\underline{k}$ are the following:

$$\mathbb{M}_{1,3} \ni \underline{x} = (x^\mu) = (x^0 = x_0, x^i = -x_i\,,\, i = 1,2,3) \equiv (x^0, \mathbf{x})\,, \tag{10}$$

equipped with the metric $\mathrm{d}s^2 = (\mathrm{d}x^0)^2 - \mathrm{d}\mathbf{x}\cdot\mathbf{x} \equiv g_{\mu\nu}\mathrm{d}x^\mu \mathrm{d}x^\nu$, $g_{\mu\nu} = \mathrm{diag}(1,-1,-1,-1)$,

$$\underline{k} = (k^\mu) = (k^0, \mathbf{k})\,. \tag{11}$$

The Minkowskian inner product is noted by:

$$\underline{x}\cdot\underline{x}' = g_{\mu\nu}x^\mu x'^\nu = x^\mu x'_\mu = x^0 x'^0 - \mathbf{x}\cdot\mathbf{x}' \tag{12}$$

Let $\underline{k}$ be a 4-momentum pointing toward point A of the mass shell hyperboloid $\mathcal{V}_m^+ = \{\underline{k}\,,\, \underline{k}\cdot\underline{k} = m^2c^2\}$, and an infinitesimal hyperbolic interval at A, with length

$$\mathrm{d}\sigma = mc\,\mathrm{d}\omega\,, \tag{13}$$

where $\mathrm{d}\omega = \dfrac{\mathrm{d}^3\mathbf{k}}{k_0}$ is the Lorentz-invariant element on $\mathcal{V}_m^+$. Given a time-like unit vector $\underline{n}$, and a straight line Δ passing through the origin and orthogonal (in the $\mathbb{M}_{1,3}$ metric sense) to $\underline{n}$, denote by $\mathrm{d}\Omega$ the length of the projection of $\mathrm{d}\sigma$ on Δ along $\underline{n}$. As is illustrated in Figure 1, one easily proves that

$$\mathrm{d}\Omega = |\underline{k}\cdot\underline{n}|\,\mathrm{d}\omega \quad (= \mathrm{d}^3\mathbf{k} \text{ if } \underline{n} = (1,\mathbf{0}))\,. \tag{14}$$

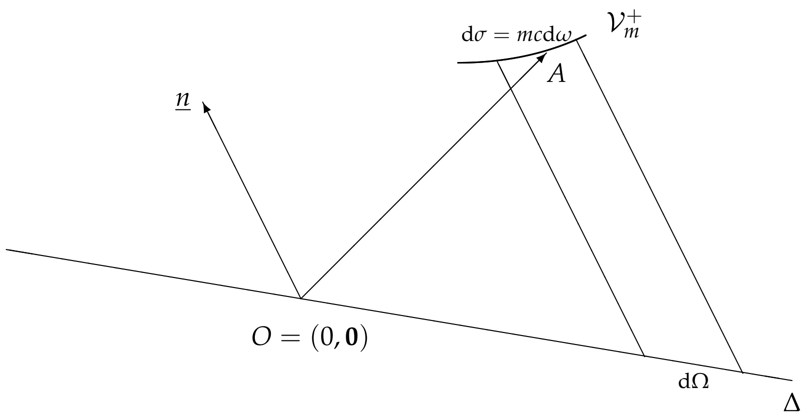

Figure 1. $\underline{n}$ is a time-like unit vector, Δ is a straight line passing through the origin and orthogonal (in the Minkowskian metric sense) to $\underline{n}$. The 4-momentum $\underline{k} = (k^\mu) = (k^0, \mathbf{k})$ points toward a point A of the mass shell hyperboloid $\mathcal{V}_m^+ = \{\underline{k}\,,\, \underline{k}\cdot\underline{k} = m^2c^2\}$. $\mathrm{d}\Omega$ is the length of the projection, along $\underline{n}$, of an infinitesimal hyperbolic interval at A of length $\mathrm{d}\sigma = mc\mathrm{d}\omega$.

The sample population consists of those particles with world lines cutting the infinitesimal space-like segment $\mathrm{d}\Sigma$ orthogonal to the time-like unit vector $\underline{n}$, as is shown in Figure 2.

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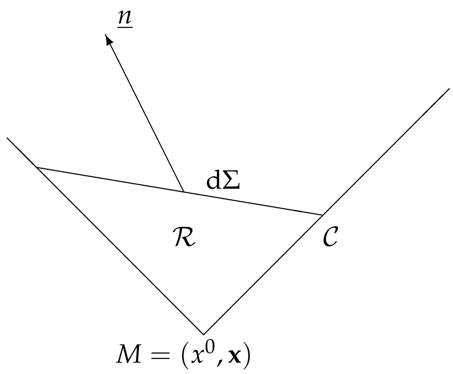

Figure 2. $\mathcal{C}$ is the portion of the null cone starting at the event $M = (x^0, \mathbf{x})$ and limited by the infinitesimal space-like segment $\mathrm{d}\Sigma$ orthogonal to the time-like unit vector $\underline{n}$. $\mathcal{R}$ is the region delimited by M, the portion of the light cone $\mathcal{C}$, and $\mathrm{d}\Sigma$.

Every particle that traverses the segment $\mathcal{C}$ of the null cone between M and $\mathrm{d}\Sigma$ must also traverse $\mathrm{d}\Sigma$ (causal cone). Consequently, regardless of the collisions that take place within the infinitesimal region $\mathcal{R}$ bounded by M, the segment of the light cone $\mathcal{C}$, and $\mathrm{d}\Sigma$, the number of particles crossing Σ is predetermined as the number crossing $\mathcal{C}$:

$$\nu = \underline{N} \cdot \underline{n}\,\mathrm{d}\Sigma = \mathrm{d}\Sigma \int_{\mathcal{V}_m^+} \mathcal{N}(\underline{x}, \underline{k})\,\mathrm{d}\Omega\,, \tag{15}$$

where $\underline{N}$ is the numerical-flux 4-vector and $\mathcal{N}(\underline{x}, \underline{k})$ is the distribution function. By the conservation of 4-momentum at each collision in a simple gas, the flux of 4-momentum across $\mathrm{d}\Sigma$ is predetermined as the flux across $\mathcal{C}$,

$$T_\mu \cdot \underline{n}\,\mathrm{d}\Sigma = \mathrm{d}\Sigma \int_{\mathcal{V}_m^+} \mathcal{N}(\underline{x}, \underline{k})\, c k_\mu \mathrm{d}\Omega\,, \tag{16}$$

where $\underline{\underline{T}} = (T_{\mu\nu})$ is the energy-momentum tensor.

The most probable distribution function $\mathcal{N}$ at M is that which maximizes the following entropy integral

$$F = -\mathrm{d}\Sigma \int_{\mathcal{V}_m^+} \mathcal{N}(\underline{x}, \underline{k})\, \log \mathcal{N}(\underline{x}, \underline{k})\,\mathrm{d}\Omega\,. \tag{17}$$

Variational calculus with 5 Lagrange $\underline{x}$-dependent multipliers α and η_μ associated with constraints on ν and $T_\mu \cdot \underline{n}$ respectively leads to the solution

$$\mathcal{N}(\underline{x}, \underline{k}) = C(\underline{x})\, \exp(-\underline{\eta}(\underline{x}) \cdot \underline{k})\,, \quad C = e^{\alpha - 1}\,. \tag{18}$$

Scalar C and time-like 4-vector $\underline{\eta}$ are determined by the constraints on $\nu = \underline{N} \cdot \underline{n}\,\mathrm{d}\Sigma$ and $T_\mu \cdot \underline{n}\,\mathrm{d}\Sigma$:

$$C \int_{\mathcal{V}_m^+} k_\mu\, e^{-\underline{\eta} \cdot \underline{k}}\,\mathrm{d}\omega = N_\mu\,, \qquad C \int_{\mathcal{V}_m^+} c k_\mu k_\nu\, e^{-\underline{\eta} \cdot \underline{k}}\,\mathrm{d}\omega = T_{\mu\nu}\,. \tag{19}$$

established by taking into account that $\underline{n}$ is arbitrary.

With the equations of conservation

$$\underline{\partial} \cdot \underline{N} = 0\,, \quad \underline{\partial} \cdot T_\mu = 0\,, \tag{20}$$

we finally get as many equations as the 19 functions of $\underline{x}$: $C, \underline{\eta}, \underline{N}, \underline{\underline{T}}$. The following partition function is essential for all relevant calculations.

$$Z(\underline{\eta}) := \int_{\mathcal{V}_m^+} e^{-\underline{\eta}\cdot\underline{k}} \frac{\mathrm{d}^3\mathbf{k}}{k_0} = \frac{4\pi mc}{\sqrt{\underline{\eta}\cdot\underline{\eta}}} K_1\left(mc\sqrt{\underline{\eta}\cdot\underline{\eta}}\right) \tag{21}$$

where K_ν is the modified Bessel function [19]. Hence, the components of the numerical flux 4-vector $\underline{N}$ and of the energy tensor $\underline{\underline{T}}$ in (19) are given in terms of derivatives of Z and finally in terms of Bessel functions by

$$N_\mu = -C\frac{\partial Z}{\partial \eta^\mu} = C\frac{4\pi m^2 c^2 \eta_\mu}{\underline{\eta}\cdot\underline{\eta}} K_2\left(mc\sqrt{\underline{\eta}\cdot\underline{\eta}}\right), \tag{22}$$

$$T_{\mu\nu} = Cc\frac{\partial^2 Z}{\partial \eta^\mu \partial \eta^\nu} = C4\pi m^2 c^3 \left[mc\frac{K_3\left(mc\sqrt{\underline{\eta}\cdot\underline{\eta}}\right)}{(\underline{\eta}\cdot\underline{\eta})^{3/2}}\eta_\mu\eta_\nu - \frac{K_2\left(mc\sqrt{\underline{\eta}\cdot\underline{\eta}}\right)}{\underline{\eta}\cdot\underline{\eta}} g_{\mu\nu}\right]. \tag{23}$$

For a simple gas consisting of material particles of proper mass m the components of the energy-momentum tensor $\underline{\underline{T}}$ are given by

$$T_{\mu\nu} = (\rho + p)u_\mu u_\nu - p g_{\mu\nu}, \tag{24}$$

where ρ is the mean density, p is the pressure, and $\underline{u} = \left(u_\mu = \frac{\mathrm{d}x_\mu}{\mathrm{d}s}\right)$, $\underline{u}\cdot\underline{u} = 1$, is the mean 4-velocity of the fluid. Hence, by identification with (23), Synge [14] proved that *a relativistic gas consisting of material particles of proper mass m is a perfect fluid* through the relations:

$$u_\mu = \frac{\eta_\mu}{\sqrt{\underline{\eta}\cdot\underline{\eta}}}, \tag{25}$$

$$\rho + p = C4\pi m^3 c^4 \frac{K_3\left(mc\sqrt{\underline{\eta}\cdot\underline{\eta}}\right)}{\sqrt{\underline{\eta}\cdot\underline{\eta}}}, \tag{26}$$

$$p = C4\pi m^2 c^3 \frac{K_2\left(mc\sqrt{\underline{\eta}\cdot\underline{\eta}}\right)}{\underline{\eta}\cdot\underline{\eta}}. \tag{27}$$

From (26) and (27) we derive the expression of the density:

$$\rho = C\frac{4\pi m^3 c^4}{\sqrt{\underline{\eta}\cdot\underline{\eta}}} \frac{K_1\left(mc\sqrt{\underline{\eta}\cdot\underline{\eta}}\right) + K_3\left(mc\sqrt{\underline{\eta}\cdot\underline{\eta}}\right)}{2} = -C\frac{4\pi m^3 c^4}{\sqrt{\underline{\eta}\cdot\underline{\eta}}} K_2'\left(mc\sqrt{\underline{\eta}\cdot\underline{\eta}}\right). \tag{28}$$

Let us define the invariant quantity, *i.e.*, the projection of the numerical flux (**??**) along the 4-velocity of the fluid,

$$\mathcal{N}_0 = \underline{N}\cdot\underline{u} = C\frac{4\pi m^2 c^2}{\sqrt{\underline{\eta}\cdot\underline{\eta}}} K_2\left(mc\sqrt{\underline{\eta}\cdot\underline{\eta}}\right). \tag{29}$$

This expression, which represents the number of particles per unit length ("numerical density") in the rest frame of the fluid ($u_0 = 1$), allows us to determine the function $C = C(\underline{x})$ and to eventually write the distribution (18) as:

$$\mathcal{N}(\underline{x}, \underline{k}) = \frac{\mathcal{N}_0}{m^2 c k_B T_a K_2\left(mc^2/k_B T_a\right)} \exp\left(-\frac{c\underline{u}\cdot\underline{k}}{k_B T_a}\right). \tag{30}$$

The term $T_a := c/(k_B\sqrt{\underline{\eta}\cdot\underline{\eta}})$, where k_B is the Boltzmann constant, is a "relativistic" absolute temperature. It is precisely the relativistic invariant, which might fit pointview (**c**).

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Note that with this expression, (27) reads as the usual gas law:

$$p = \mathcal{N}_0 k_B T_a \,. \tag{31}$$

The Maxwell-Boltzmann non relativistic distribution (in the space of momenta) is recovered by considering the limit at $k_B T_a \ll mc^2$ in the rest frame of the fluid:

$$\begin{aligned} & K_2\left(\frac{mc^2}{k_B T_a}\right) \approx \sqrt{\frac{\pi k_B T_a}{2mc^2}} e^{-\frac{mc^2}{k_B T_a}} \\ & \Rightarrow \mathcal{N}(\underline{x}, \underline{k}) \\ & \approx N_0 (2\pi m k_B T_a)^{-3/2} \exp\left(-\frac{k_0 c - mc^2}{k_B T_a}\right) \approx N_0 (2\pi m k_B T_a)^{-3/2} \exp\left(-\frac{\mathbf{k}^2}{2m k_B T_a}\right) . \end{aligned} \tag{32}$$

3.1. Inverse Temperature 4-Vector

The found distribution (30) on the Minkowskian mass shell for a simple gas consisting of particles of proper mass m leads us to introduce the relativistic thermodynamic, future directed, time-like 4-coldness vector $\underline{\beta}$, as the 4-version of the reciprocal of the thermodynamic temperature (see also [2]):

$$\frac{c\underline{u}}{k_B T_a} \equiv \underline{\beta} = (\beta^0 = \beta_0 > 0, \beta^i = -\beta_i) = (\beta_0, \boldsymbol{\beta}) \,, \tag{33}$$

with *absolute coldness* as relativistic invariant,

$$\sqrt{\underline{\beta} \cdot \underline{\beta}} = \frac{c}{k_B T_a} \equiv \beta_a \,. \tag{34}$$

It is precisely the way as the component β_0 transforms under a Lorentz boost, $\beta_0' = \gamma(v)(\beta_0 - \mathbf{v} \cdot \boldsymbol{\beta}/c)$, which explains the way the temperature transforms à la de Broglie, $T \mapsto T' = T\gamma^{-1}$. So, in the sequel, we call Maxwell-Jüttner distribution the following relativistic invariant:

$$\mathcal{N}(\underline{\beta}, \underline{k}) = \frac{\mathcal{N}_0}{mcK_1\left(mc\beta_a\right)} \exp\left(-\underline{\beta} \cdot \underline{k}\right) \,, \tag{35}$$

where the space-time dependence holds through the coldness 4-vector coldness field $\underline{\beta} = \underline{\beta}(\underline{x})$.

4. de Sitter Material

We now turn our attention to the de Sitter (dS) space-time and some important features of a dS covariant quantum field theory.

4.1. de Sitter Geometry

The de Sitter space-tiem can be viewed as a hyperboloid embedded in a five-dimensional Minkowski space $\mathbb{M}_{1,4}$ with metric $g_{\alpha\beta}$ =diag$(1, -1, -1, -1, -1)$ (see Figure 3). Of course, one should keep in mind that all choices of one point in the manifold as an origin are physically equivalent, like are the points of the Minkowski space-time $\mathbb{M}_{1,3}$.

$$\mathrm{M}_R \equiv \{x \in \mathbb{R}^5;\ x^2 = g_{\alpha\beta}\, x^\alpha x^\beta = -R^2\}, \,, \quad \alpha, \beta = 0,1,2,3,4 \,, \tag{36}$$

where the pseudo-radius R (or inverse of curvature) is given by $R = \sqrt{\frac{3}{\Lambda}}$ within the cosmological ΛCDM standard model. The de Sitter symmetry group is the group $\mathrm{SO}_0(1,4)$ of proper (i.e., det$\,. = 1$) and orthochronous (to be precised later) transformations of the manifold (36). This group has ten (Killing) generators $K_{\alpha\beta} = x_\alpha \partial_\beta - x_\beta \partial_\alpha$.

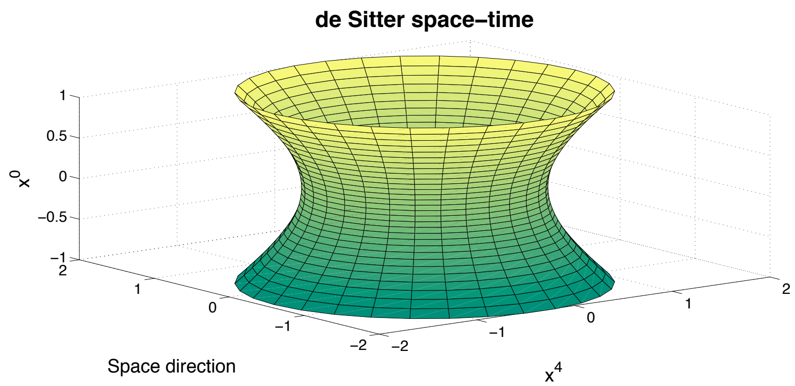

Figure 3. The de Sitter space-time as viewed as a one-sheet hyperboloid embedded in Minkowski space $\mathbb{M}_{1,4}$.

4.2. Flat Minkowskian Limit of de Sitter Geometry

Let us choose the global coordinates $ct \in \mathbb{R}, \mathbf{n} \in \mathbb{S}^2, r/R \in [0,\pi]$ for the dS manifold M_R. They are defined by:

$$\begin{aligned}\mathrm{M}_R \ni x &= (x^0, x^1, x^2, x^3, x^4) \equiv (x^0, \mathbf{x}, x^4) \\ &= (R\sinh(ct/R),\ R\cosh(ct/R)\sin(r/R)\mathbf{n}, R\cosh(ct/R)\cos(r/R)) \equiv x(t,\mathbf{x})\,. \end{aligned} \tag{37}$$

At the limit $R \to \infty$ the manifold $\mathrm{M}_R \to \mathbb{M}_{1,3}$, the Minkowski spacetime tangent to M_R at, say, the de Sitter point $O_{dS} = (0, \mathbf{0}, R)$, chosen as origin, since then

$$\mathrm{M}_R \ni x \underset{R\to\infty}{\approx} (ct, \mathbf{r} = r\,\mathbf{n}, R) \equiv (\underline{\ell}, R)\,, \quad \underline{\ell} \in \mathbb{M}_{1,3}\,. \tag{38}$$

At this limit, the de Sitter group becomes the Poincaré group:

$$\lim_{R\to\infty} \mathrm{SO}_0(1,4) = \mathcal{P}_+^{\uparrow}(1,3) = \mathbb{M}_{1,3} \rtimes \mathrm{SO}_0(1,3)\,. \tag{39}$$

Consistently, the ten de Sitter Killing generators contract (in the Wigner-Inonü sense) to their Poincaré counterparts $K_{\mu\nu}, \Pi_\mu, \mu = 0,1,2,3$, after rescaling the four $K_{4\mu} \longrightarrow \Pi_\mu = K_{4\mu}/R$.

4.3. de Sitter plane waves as binomial deformations of Minkowskian plane waves

The de Sitter (scalar) plane waves are defined [20] as

$$\phi_{\tau,\xi}(x) = \left(\frac{x\cdot\xi}{R}\right)^{\tau}, \quad x \in \mathrm{M}_R\,, \quad \xi \in \mathcal{C}_{1,4}\,, \tag{40}$$

where $\mathcal{C}_{1,4} = \{\xi \in \mathbb{R}^5\,,\, \xi\cdot\xi = 0\}$ is the null cone in $\mathbb{M}_{1,4}$. They are solutions of the Klein-Gordon-like equation

$$\frac{1}{2}M_{\alpha\beta}M^{\alpha\beta}\phi_{\tau,\xi}(x) \equiv R^2\Box_R\phi_{\tau,\xi}(x) = \tau(\tau+3)\phi_{\tau,\xi}(x)\,,$$

where $M_{\alpha\beta} = -\mathrm{i}\left(x_\alpha\partial_\beta - x_\beta\partial_\alpha\right)$ is the quantum representation of the Killing vector $K_{\alpha\beta}$, and $\Box_R$ stands for the d'Alembertian operator on M_R. For the values

$$\tau = -\frac{3}{2} + \mathrm{i}\nu\,, \quad \nu \in \mathbb{R}\,, \tag{41}$$

they describe free quantum motions of "massive" scalar particles on M_R. The term "massive" is justified by the flat Minkowskian limit $R \to \infty$, i.e. $\Lambda \to 0$. This limit is understood as follows.

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(**i**) First one has the Garidi [21] relation between proper mass m (curvature independent) of the spinless particle and the parameter $\nu \geq 0$:

$$m = \frac{\hbar}{Rc}\left[\nu^2 + \frac{1}{4}\right]^{1/2} \Leftrightarrow \nu = \sqrt{\frac{R^2m^2c^2}{\hbar^2} - \frac{1}{4}} \underset{R\,\text{large}}{\approx} \frac{Rmc}{\hbar} = \frac{mc}{\hbar}\sqrt{\frac{3}{\Lambda}}\,. \tag{42}$$

The quantity $\dfrac{\hbar c\nu}{R}$ is a kind of *at rest desitterian energy*, which is distinct of the proper mass energy mc^2 if $\Lambda \neq 0$.

(**ii**) Then, with the mass shell parametrisation $\xi = \left(\xi^0 = \frac{k_0}{mc}, \boldsymbol{\xi} = \frac{\mathbf{k}}{mc}, \xi^4 = 1\right) \in \mathcal{C}^+_{1,4}$, one obtains at the limit $R \to \infty$:

$$\phi_{\tau,\xi}(x) = (x\cdot\xi/R)^{-3/2+\mathrm{i}\nu} \underset{R\to\infty}{\to} e^{\mathrm{i}k\cdot\underline{\ell}/\hbar}\,, \quad \underline{\ell} = (ct, \mathbf{r})\,. \tag{43}$$

This relation allows to consider the functions (40) as deformation of the plane waves propagating in the Minkowskian space-time $\mathbb{M}_{1,4}$. This pivotal property justifies the name "dS plane waves" granted to the functions (40).

4.4. Analytic Extension of dS Plane Waves for dS QFT

The dS plane waves $\phi_{\tau,\xi}(x) = \left(\dfrac{x\cdot\xi}{R}\right)^{\tau}$, $\tau = -3/2 + \mathrm{i}\nu$, are not defined on all M_R due to the possible change of sign of $x\cdot\xi$. A solution to this drawback is found through the extension to the tubular domains in the complexified hyperboloid $\mathrm{M}_R^{\mathbb{C}} = \{z = x + \mathrm{i}y \in \mathbb{C}^5\,,\, z^2 = g_{\alpha\beta}\, z^\alpha z^\beta = -R^2$ or equivalently $x^2 - y^2 = -R^2\,,\, x\cdot y = 0\}$:

$$\mathcal{T}^{\pm} := \mathrm{T}^{\pm} \cap \mathrm{M}_R^{\mathbb{C}}\,, \quad \mathrm{T}^{\pm} := \mathbb{M}_{1,4} + \mathrm{i}\mathrm{V}^{\pm}\,, \tag{44}$$

where the forward and backward light cones $\mathrm{V}^{\pm} := \left\{x \in \mathbb{M}_{1,4}\,,\, x^0 \gtrless \sqrt{\mathbf{x}^2 + (x^4)^2}\right\}$ allow for a causal ordering in $\mathbb{M}_{1,4}$.

Then the extended plane waves $\phi_{\tau,\xi}(z) = \left(\dfrac{z\cdot\xi}{R}\right)^{\tau}$ are globally defined for $z \in \mathcal{T}^{\pm}$ and $\xi \in \mathcal{C}^+_{1,4}$.

These analytic extensions allow for a consistent QFT for free scalar fields on M_R: the two-point Wightman function $\mathcal{W}_\nu(x,x') = \langle\Omega, \phi(x)\phi(x')\Omega\rangle$ can be extended to the complex covariant, maximally analytic, two-point function having the spectral representation in terms of these extended plane waves:

$$W_\nu(z,z') = c_\nu \int_{\mathcal{V}_m^+\cup\mathcal{V}_m^-} (z\cdot\xi)^{-3/2+\mathrm{i}\nu}(\xi\cdot z')^{-3/2-\mathrm{i}\nu}\,\frac{\mathrm{d}\mathbf{k}}{k_0}\,, \quad z \in \mathcal{T}^-\,,\, z' \in \mathcal{T}^+\,. \tag{45}$$

Details are found in [20] and in the recent volume [22].

4.5. KMS Interpretation of $W_\nu(z,z')$ Analyticity

From the analyticity of $W_\nu(z,z')$ we deduce that $\mathcal{W}_\nu(x,x')$ defines a $2\mathrm{i}\pi R/c$ periodic analytic function of t, whose domain is the periodic cut plane

$$\mathbb{C}^{\mathrm{cut}}_{x,x'} = \{t \in \mathbb{C}\,,\, \mathrm{Im}(t) \neq 2n\pi R/c\,,\, n \in \mathbb{Z}\} \cup \{t\,,\, t - 2\mathrm{i}n\pi R/c \in I_{x,x'}\,,\, n \in \mathbb{Z}\}\,, \tag{46}$$

where $I_{x,x'}$ is the real interval on which $(x-x')^2 < 0$. Hence $W_\nu(z,z')$ is analytic in the strip

$$\{t \in \mathbb{C}\,,\, 0 < \mathrm{Im}(t) < 2\mathrm{i}\pi R/c\}\,, \tag{47}$$

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and satisfies:

$$\mathcal{W}_\nu(x'(t+t',\mathbf{x}),x) = \lim_{\epsilon\to 0^+} W_\nu\left((x, x'(t+t'+2\mathrm{i}\pi R/c - \mathrm{i}\epsilon, \mathbf{x})\right)\,, \quad t' \in \mathbb{R}\,. \tag{48}$$

This is a KMS relation at ($\sim$ Hawking) temperature

$$T_\Lambda = \frac{\hbar c}{2\pi k_B R} := \frac{\hbar c}{2\pi k_B}\sqrt{\frac{\Lambda}{3}}\,. \tag{49}$$

5. de Sitterian Tsallis Distribution

5.1. Tsallis Entropy and Distribution, a Short Reminder

Given a discrete (resp. continuous) set of probabilities $\{p_i\}$ (resp. continuous $x \mapsto p(x)$) with $\sum_i p_i = 1$ (resp. $\int p(x)\mathrm{d}x = 1$) , and a real q, the Tsallis entropy [23] is defined as

$$S_q(p_i) = k\frac{1}{q-1}\left(1 - \sum_i p_i^q\right) \quad \text{resp.} \quad S_q[p] = \frac{1}{q-1}\left(1 - \int (p(x))^q \mathrm{d}x\right)\,. \tag{50}$$

As $q \to 1$, $S_q(p_i) \to S_{\mathrm{BG}}(p) = -k\sum_i p_i \ln p_i$ (Boltzmann-Gibbs). The Tsallis entropy is non additive : for two independent systems A and B, for which $p(A\cup B) = p(A)\,p(B)$, $S_q(A\cup B) = S_q(A) + S_q(B) + (1-q)S_q(A)S_q(B)$. A *Tsallis distribution* is a probability distribution derived from the maximization of the Tsallis entropy under appropriate constraints. The so-called q-exponential Tsallis distribution has the probability density function

$$(2-q)\lambda[1-(1-q)\lambda x]^{1/(1-q)} \equiv (2-q)\lambda e_q(-\lambda x)\,, \tag{51}$$

where $q < 2$ and $\lambda > 0$ (rate), arises from the maximization of the Tsallis entropy under appropriate constraints, including constraining the domain to be positive. More details are given for instance in [24].

Let us now show how the Tsallis distribution can be viewed as a Λ-deformation of the Maxwell-Jüttner distribution.

5.2. Coldness in de Sitter

In analogy with the de Sitter plane waves, we introduce the following distributions on the subset $\sim \mathcal{V}_m^+$ of the null cone $\mathcal{C}_{1,4}^+ = \{\xi \in \mathbb{M}_{1,4}\,,\, \xi\cdot\xi = 0\,,\, \xi^0 > 0\}$:

$$\phi_{\tau,\xi}(x) = \left(\frac{\mathfrak{b}\cdot\xi}{B}\right)^\tau\,, \quad \mathfrak{b} \in \mathsf{M}_B\,, \quad \xi = \left(\frac{k^0}{mc} > 0, \frac{\mathbf{k}}{mc}, -1\right)\,, \tag{52}$$

where one should note the negative value -1 for ξ_4, and

$$\mathsf{M}_B \equiv \{\mathfrak{b} \in \mathbb{M}_{1,4}\,,\, \mathfrak{b}^2 = g_{\alpha\beta}\,\mathfrak{b}^\alpha\mathfrak{b}^\beta = -B^2\}, \quad \alpha,\beta = 0,1,2,3,4\,, \tag{53}$$

is the manifold of the “deSitterian 5-vector coldness fields” $\mathfrak{b} = \mathfrak{b}(x)$.

Like for M_R we use global coordinates on M_B:

$$\beta^0 \in \mathbb{R}\,, \quad \boldsymbol{\beta} = \|\boldsymbol{\beta}\|\mathbf{n} \in \mathbb{R}^3\,, \quad \|\boldsymbol{\beta}\|/B \in [0,\pi]\,, \tag{54}$$

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with

$$\begin{aligned} \mathsf{M}_B \ni \mathfrak{b} &\equiv \mathfrak{b}(\underline{\beta}) = (\mathfrak{b}^0, \mathfrak{b}^1, \mathfrak{b}^2, \mathfrak{b}^3, \mathfrak{b}^4) \equiv (\mathfrak{b}^0, \boldsymbol{\mathfrak{b}}, \mathfrak{b}^4) \\ &= \Big(B\sinh(\beta^0/B),\, B\cosh(\beta^0/B)\sin(\|\boldsymbol{\beta}\|/B)\mathbf{n}, -B\cosh(\beta^0/B)\cos(\|\boldsymbol{\beta}\|/B)\Big)\,, \end{aligned} \tag{55}$$

in such a way that at large B we recover the Minkowskian coldness $\underline{\beta}$:

$$\mathsf{M}_B \ni \mathfrak{b} \underset{B\to\infty}{\sim} (\underline{\beta}, B)\,.$$

We now need to connect the desitterian coldness scale B with Λ. Inspired by the relativistic invariant $\beta_a = \dfrac{c}{k_B T_a}$ and the KMS temperature $T_\Lambda = \frac{\hbar c}{2\pi k_B}\sqrt{\frac{\Lambda}{3}}$ we write

$$B \propto \frac{2\pi}{\hbar}\sqrt{\frac{3}{\Lambda}}\,, \quad \text{i.e.} \quad B = \frac{\mathfrak{n}}{\hbar\sqrt{\Lambda}}\,, \tag{56}$$

where $\mathfrak{n}$ is a numerical factor. Note that with the values

$$\Lambda_{\text{current}} = 1.1056 \times 10^{-52}\,\mathrm{m}^{-2}\,, \quad \hbar = 1.054571817... \times 10^{-34}\,\mathrm{J\,s}\,,$$

one obtains $B \approx 0.9 \times 10^{60}\,\mathfrak{n}\,\mathrm{SI}$ (inverse of a momentum).

5.3. A de Sitterian Tsallis Distribution

We now consider the distribution on $\mathsf{M}_B \times \mathcal{V}_m^+$ with $B = \frac{\mathfrak{n}}{\hbar\sqrt{\Lambda}}$:

$$\begin{gathered} \mathcal{N}(\mathfrak{b}, \underline{k}) = C_B \left(\frac{\mathfrak{b}\cdot\xi}{B}\right)^{-mcB} = C_B\left(\frac{\mathfrak{b}^0}{B}\frac{k^0}{mc} - \frac{\boldsymbol{\mathfrak{b}}}{B}\cdot\frac{\mathbf{k}}{mc} + \frac{\mathfrak{b}^4}{B}\right)^{-mcB}\,. \\ \mathfrak{b} \in \mathsf{M}_B\,, \quad \xi = \left(\frac{k^0}{mc} > 0, \frac{\mathbf{k}}{mc}, -1\right)\,, \end{gathered} \tag{57}$$

where the constant C_B involves an associated Legendre function of the First Kind [25].

With the global coordinates (55) and with the constraint $\beta^0/B \in [0, \pi/2)$, the distribution $\mathcal{N}(\mathfrak{b}, \underline{k})$ reads

$$\begin{aligned} &\mathcal{N}(\mathfrak{b}, \underline{k}) \\ &= C_B\left(\cosh(\beta^0/B)\cos(\|\boldsymbol{\beta}\|/B) + \sinh(\beta^0/B)\frac{k^0}{mc} - \cosh(\beta^0/B)\sin(\|\boldsymbol{\beta}\|/B)\frac{\mathbf{n}\cdot\mathbf{k}}{mc}\right)^{-mcB} \\ &= C_B \exp\left[-mcB\log\left(\cosh(\beta^0/B)\cos(\|\boldsymbol{\beta}\|/B)\right)\right] \times \\ &\times \exp\left[-mcB\log\left(1 + \frac{\sinh(\beta^0/B)\frac{k^0}{mc} - \cosh(\beta^0/B)\sin(\|\boldsymbol{\beta}\|/B)\frac{\mathbf{n}\cdot\mathbf{k}}{mc}}{\cosh(\beta^0/B)\cos(\|\boldsymbol{\beta}\|/B)}\right)\right]\,. \end{aligned} \tag{58}$$

At large B this expression becomes the Maxwell-Jüttner distribution:

$$\mathcal{N}(\mathfrak{b}, \underline{k}) \underset{B\to\infty}{\sim} C_B e^{-\underline{\beta}\cdot\underline{k}}\,.$$

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Hence, going back to the original expression

$$\begin{aligned}\mathcal{N}(\mathfrak{b},\underline{k}) &= C_B\left(\frac{\mathfrak{b}\cdot\xi}{B}\right)^{-mcB} = C_B\left(\frac{\mathfrak{b}^0}{B}\frac{k^0}{mc} - \frac{\boldsymbol{\mathfrak{b}}}{B}\cdot\frac{\mathbf{k}}{mc} + \frac{\mathfrak{b}^4}{B}\right)^{-mcB}\\ &= C_B\left(\frac{\mathfrak{b}^4}{B}\right)^{-mcB}\left(1+\frac{\underline{\mathfrak{b}}\cdot\underline{k}}{\mathfrak{b}^4 mc}\right)^{-mcB}, \quad \underline{\mathfrak{b}} := (\mathfrak{b}^0, \boldsymbol{\mathfrak{b}}),\end{aligned}$$

and introducing

$$q = 1 + \frac{1}{mcB} = 1 + \frac{\hbar\sqrt{\Lambda}}{mc\mathfrak{n}}, \tag{59}$$

we finally get the Tsallis-type distribution

$$\mathcal{N}(\mathfrak{b},\underline{k}) = C_B\left(\frac{\mathfrak{b}^4}{B}\right)^{-mcB}\left(1-(1-q)\frac{B}{\mathfrak{b}^4}\,\underline{\mathfrak{b}}\cdot\underline{k}\right)^{\frac{1}{1-q}}. \tag{60}$$

Analogously to (21) and all subsequent determinations of thermodynamical quantities, the following partition function is essential for their transcriptions to the de Sitter case:

$$Z_{\mathrm{dS}}(\mathfrak{b},\underline{k}) = \left(\frac{\mathfrak{b}^4}{B}\right)^{-mcB}\int_{\mathcal{V}_m^+}\left(1+\frac{\underline{\mathfrak{b}}\cdot\underline{k}}{\mathfrak{b}^4 mc}\right)^{-mcB}\frac{\mathrm{d}^3\mathbf{k}}{k_0} \tag{61}$$

$$= 4\pi m^2c^2\left(\frac{\mathfrak{b}^4}{B}\right)^{-mcB}\int_0^\infty\left(1+\left(\frac{\mathfrak{b}_0}{\mathfrak{b}^4}\right)\cosh t\right)^{-mcB}\sinh^2 t\,\mathrm{d}t\,. \tag{62}$$

With the following integral representation of the associated Legendre function of the First Kind $P_\nu^\mu(z)$ [25],

$$P_\nu^\mu(z) = \frac{2^{-\nu}\,(z^2-1)^{-\mu/2}}{\Gamma(-\nu-\mu)\Gamma(\nu+1)}\int_0^\infty (z+\cosh t)^{-\nu-\mu-1}\sinh^{2\nu+1} t\,\mathrm{d}t\,, \tag{63}$$

valid for $z \notin (-\infty,-1]$ and $\mathrm{Re}(-\mu) > \mathrm{Re}(\nu) > -1$, the function (61) reads as

$$Z_{\mathrm{dS}}(\mathfrak{b},\underline{k}) = (8\pi)^{3/2}\Gamma(1-mcB)\left(\frac{B}{\mathfrak{b}_0}\right)^{mcB}\left(\frac{B^2-\boldsymbol{\mathfrak{b}}\cdot\boldsymbol{\mathfrak{b}}}{\mathfrak{b}_0^2}\right)^{mcB/2-3/4} P_{1/2}^{mcB-3/2}\left(\frac{\mathfrak{b}^4}{\mathfrak{b}_0}\right). \tag{64}$$

6. Conclusions

In this contribution, we have forged a groundbreaking link between the Tsallis distribution, quantum statistics, and the cosmological constant, illuminating the complex interplay between relativistic thermodynamics and a fundamental cosmological parameter.

Our key findings are encapsulated in the formulas (59) and (60). The intricate technical details of the associated thermodynamic features (flux number, energy-momentum tensor, etc.) in the de Sitter space-time, along with their physical (and astrophysical!) implications, are reserved for future exploration.

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